



Star Formation Laws and Efficiencies across 80 Nearby Galaxies

Jiayi Sun (孙嘉懿)^{1,2,39}, Adam K. Leroy^{3,4}, Eve C. Ostriker⁵, Sharon Meidt⁶, Erik Rosolowsky⁷, Eva Schinnerer⁸,
 Christine D. Wilson¹, Dyas Utomo⁹, Francesco Belfiore¹⁰, Guillermo A. Blanc^{11,12}, Eric Emsellem^{13,14},
 Christopher Faesi¹⁵, Brent Groves¹⁶, Annie Hughes¹⁷, Eric W. Koch¹⁸, Kathryn Kreckel¹⁹, Daizhong Liu²⁰,
 Hsi-An Pan²¹, Jérôme Pety^{22,23}, Miguel Querejeta²⁴, Alessandro Razza¹², Toshiki Saito²⁵, Amy Sardone^{3,4},
 Antonio Usero²⁴, Thomas G. Williams^{8,26}, Frank Bigiel²⁷, Alberto D. Bolatto²⁸, Mélanie Chevance^{29,30},
 Daniel A. Dale³¹, Jindra Gensior³², Simon C. O. Glover²⁹, Kathryn Grasha^{33,34}, Jonathan D. Henshaw^{8,35},
 María J. Jiménez-Donaire²⁴, Ralf S. Klessen^{29,36}, J. M. Diederik Kruijssen³⁰, Eric J. Murphy⁹, Lukas Neumann²⁷,
 Yu-Hsuan Teng³⁷, and David A. Thilker³⁸

¹ Department of Physics and Astronomy, McMaster University, 1280 Main Street West, Hamilton, ON L8S 4M1, Canada; sun208@mcmaster.ca

² Canadian Institute for Theoretical Astrophysics (CITA), University of Toronto, 60 St George Street, Toronto, ON M5S 3H8, Canada

³ Department of Astronomy, The Ohio State University, 140 West 18th Avenue, Columbus, OH 43210, USA

⁴ Center for Cosmology and Astroparticle Physics (CCAPP), 191 West Woodruff Avenue, Columbus, OH 43210, USA

⁵ Department of Astrophysical Sciences, Princeton University, Princeton, NJ 08544 USA

⁶ Sterrenkundig Observatorium, Universiteit Gent, Krijgslaan 281 S9, B-9000 Gent, Belgium

⁷ Department of Physics, University of Alberta, Edmonton, AB T6G 2E1, Canada

⁸ Max-Planck-Institut für Astronomie, Königstuhl 17, D-69117, Heidelberg, Germany

⁹ National Radio Astronomy Observatory, 520 Edgemont Road, Charlottesville, VA 22903-2475, USA

¹⁰ INAF—Osservatorio Astrofisico di Arcetri, Largo E. Fermi 5, I-50157, Firenze, Italy

¹¹ Observatories of the Carnegie Institution for Science, 813 Santa Barbara Street, Pasadena, CA 91101, USA

¹² Departamento de Astronomía, Universidad de Chile, Camino del Observatorio 1515, Las Condes, Santiago, Chile

¹³ European Southern Observatory, Karl-Schwarzschild Straße 2, D-85748 Garching bei München, Germany

¹⁴ Univ Lyon, Univ Lyon 1, ENS de Lyon, CNRS, Centre de Recherche Astrophysique de Lyon UMR5574, F-69230 Saint-Genis-Laval, France

¹⁵ University of Massachusetts Amherst, 710 N. Pleasant Street, Amherst, MA 01003, USA

¹⁶ International Centre for Radio Astronomy Research, University of Western Australia, 35 Stirling Highway, Crawley, WA 6009, Australia

¹⁷ CNRS, IRAP, 9 Av. du Colonel Roche, BP 44346, F-31028 Toulouse cedex 4, France

¹⁸ Center for Astrophysics | Harvard & Smithsonian, 60 Garden Street, Cambridge, MA 02138, USA

¹⁹ Astronomisches Rechen-Institut, Zentrum für Astronomie der Universität Heidelberg, Mönchhofstraße 12-14, D-69120 Heidelberg, Germany

²⁰ Max-Planck-Institut für extraterrestrische Physik, Giessenbachstraße 1, D-85748 Garching, Germany

²¹ Department of Physics, Tamkang University, No.151, Yingzhuang Rd., Tamsui Dist., New Taipei City 251301, Taiwan

²² Institut de Radioastronomie Millimétrique (IRAM), 300 Rue de la Piscine, F-38406 Saint Martin d'Hères, France

²³ Sorbonne Université, Observatoire de Paris, Université PSL, CNRS, LERMA, F-75014, Paris, France

²⁴ Observatorio Astronómico Nacional (IGN), C/Alfonso XII, 3, E-28014 Madrid, Spain

²⁵ National Astronomical Observatory of Japan, 2-21-1 Osawa, Mitaka, Tokyo, 181-8588, Japan

²⁶ Sub-department of Astrophysics, Department of Physics, University of Oxford, Keble Road, Oxford OX1 3RH, UK

²⁷ Argelander-Institut für Astronomie, Universität Bonn, Auf dem Hügel 71, D-53121 Bonn, Germany

²⁸ Department of Astronomy and Joint Space-Science Institute, University of Maryland, College Park, MD 20742, USA

²⁹ Universität Heidelberg, Zentrum für Astronomie, Institut für Theoretische Astrophysik, Albert-Ueberle-Str 2, D-69120 Heidelberg, Germany

³⁰ Cosmic Origins Of Life (COOL) Research DAO⁴⁰

³¹ Department of Physics and Astronomy, University of Wyoming, Laramie, WY 82071, USA

³² Institute for Computational Science, University of Zürich, Winterthurerstrasse 190, 8057 Zürich, Switzerland

³³ Research School of Astronomy and Astrophysics, Australian National University, Canberra, ACT 2611, Australia

³⁴ ARC Centre of Excellence for All Sky Astrophysics in 3 Dimensions (ASTRO 3D), Australia

³⁵ Astrophysics Research Institute, Liverpool John Moores University, IC2, Liverpool Science Park, 146 Brownlow Hill, Liverpool L3 5RF, UK

³⁶ Universität Heidelberg, Interdisziplinäres Zentrum für Wissenschaftliches Rechnen, Im Neuenheimer Feld 205, D-69120 Heidelberg, Germany

³⁷ Center for Astrophysics and Space Sciences, Department of Physics, University of California, San Diego, 9500 Gilman Drive, La Jolla, CA 92093, USA

³⁸ Department of Physics and Astronomy, The Johns Hopkins University, Baltimore, MD 21218, USA

Received 2022 November 17; revised 2023 February 17; accepted 2023 February 19; published 2023 March 8

Abstract

We measure empirical relationships between the local star formation rate (SFR) and properties of the star-forming molecular gas on 1.5 kpc scales across 80 nearby galaxies. These relationships, commonly referred to as “star formation laws,” aim at predicting the local SFR surface density from various combinations of molecular gas surface density, galactic orbital time, molecular cloud free fall time, and the interstellar medium dynamical equilibrium pressure. Leveraging a multiwavelength database built for the Physics at High Angular Resolution in Nearby Galaxies (PHANGS) survey, we measure these quantities consistently across all galaxies and quantify systematic uncertainties stemming from choices of SFR calibrations and the CO-to-H₂ conversion factors. The star formation laws we examine show 0.3–0.4 dex of intrinsic scatter, among which the molecular Kennicutt–Schmidt

³⁹ CITA National Fellow.

⁴⁰ coolresearch.io



Original content from this work may be used under the terms of the [Creative Commons Attribution 4.0 licence](https://creativecommons.org/licenses/by/4.0/). Any further distribution of this work must maintain attribution to the author(s) and the title of the work, journal citation and DOI.

relation shows a $\sim 10\%$ larger scatter than the other three. The slope of this relation ranges $\beta \approx 0.9\text{--}1.2$, implying that the molecular gas depletion time remains roughly constant across the environments probed in our sample. The other relations have shallower slopes ($\beta \approx 0.6\text{--}1.0$), suggesting that the star formation efficiency per orbital time, the star formation efficiency per free fall time, and the pressure-to-SFR surface density ratio (i.e., the feedback yield) vary systematically with local molecular gas and SFR surface densities. Last but not least, the shapes of the star formation laws depend sensitively on methodological choices. Different choices of SFR calibrations can introduce systematic uncertainties of at least $10\%\text{--}15\%$ in the star formation law slopes and $0.15\text{--}0.25$ dex in their normalization, while the CO-to-H₂ conversion factors can additionally produce uncertainties of $20\%\text{--}25\%$ for the slope and $0.10\text{--}0.20$ dex for the normalization.

Unified Astronomy Thesaurus concepts: [Interstellar medium \(847\)](#); [Star formation \(1569\)](#); [Galaxy evolution \(594\)](#); [Scaling relations \(2031\)](#)

Supporting material: figure set, machine-readable table

1. Introduction

“Star formation laws” are empirical scaling relations between properties of the interstellar gas and the star formation rate (SFR) of this gas. These relations arise from the physical processes governing star formation (SF) in the interstellar medium (ISM) in galaxies near and far (see review by Kennicutt & Evans 2012).

Since the pioneering work of Schmidt (1959), many forms of SF laws have been proposed in the literature. One large family is known as “integrated” SF laws, which connect unresolved, global measurements of galaxy gas mass and SFR (e.g., Kennicutt 1998; Saintonge et al. 2011; de los Reyes & Kennicutt 2019). The other family describes “resolved” SF laws, which relate the local surface densities of gas mass and SFR (usually measured at \sim kiloparsec scales; e.g., Wong & Blitz 2002; Bigiel et al. 2008). Alternative formulations have modified the basic relationship by, for example, (a) including only molecular gas (e.g., Wong & Blitz 2002; Bigiel et al. 2011) or dense molecular gas (e.g., Gao & Solomon 2004; Lada et al. 2012) as opposed to the total neutral gas, (b) considering volume densities instead of surface densities (e.g., Schmidt 1959; Bacchini et al. 2019), or (c) incorporating additional information beyond gas mass/densities for the independent variable (such as orbital time; see Elmegreen 1997; Silk 1997). The functional forms of some of the resolved empirical SF laws are originally motivated by theoretical considerations, as we shall discuss below.

Among the resolved SF laws, at least four of them have attracted greater attention in recent decades. These are (1) the molecular Kennicutt–Schmidt (mKS) relation (Kennicutt 1998) between the surface densities of molecular gas and SFR; (2) the molecular Elmegreen–Silk (mES) relation (Elmegreen 1997; Silk 1997) between the SFR surface density and molecular gas surface density divided by orbital time; (3) the free fall time-regulated SF relation (FFTR relation; McKee & Ostriker 2007; Krumholz et al. 2009, 2012) linking the SFR surface density to the molecular gas surface density divided by the molecular cloud free fall time; and (4) the pressure-regulated SF relation (PR relation; Ostriker et al. 2010; Ostriker & Shetty 2011) connecting the SFR surface density to the ISM dynamical equilibrium pressure. Many works have suggested near-unity slopes for these relations (between 0.8 and 1.2; see, e.g., Daddi et al. 2010; Genzel et al. 2010; Bigiel et al. 2011; Krumholz et al. 2012; Ostriker & Kim 2022), which implies that the ratio of the dependent and independent variables (i.e., the proportionality constant) remains roughly unchanged across a wide range of physical conditions. These SF laws and the

corresponding proportionality constants (namely, the molecular gas depletion time, the SF efficiency per unit orbital time and per unit free fall time, and the feedback yield) are thus of great interest and have been linked to various SF theories.

In this Letter, we intend to accomplish three overarching goals. First, we aim to provide the latest measurements of these four SF laws and their associated proportionality constants across 80 nearby, star-forming galaxies mapped by the Physics at High Angular resolution in Nearby Galaxies with Atacama Large Millimeter/submillimeter Array survey (PHANGS–ALMA; Leroy et al. 2021a, 2021b). The excellent depth, resolution, and field-of-view coverage of the PHANGS–ALMA survey allow us to measure these relations throughout a representative range of star-forming environments in the local universe. Second, we compare these SF laws in terms of their overall slopes and scatter, as well as the scatter in the corresponding proportionality constants. These measurements provide an empirical basis for determining the predictive power of these SF laws for the local SFR. Third, we examine how the best-fit slope, normalization, and scatter of each SF law depends on the approach used to estimate physical quantities, especially the SFR and molecular gas mass, from observable quantities. Such estimation often relies on various methodological choices (such as a particular SFR calibration or CO-to-H₂ conversion factor), and many different choices have been proposed for different physical regimes or under different observational limitations. Quantitative comparisons among even a subset of these choices can provide useful estimates for the systematic uncertainties they introduce.

We note that the mKS relation has been reported for PHANGS galaxies in various subsamples and subgalactic environments (see Pessa et al. 2021; Querejeta et al. 2021); the SF efficiency per free fall time has been measured within the PHANGS–ALMA pilot sample (Utomo et al. 2018); and the PR relation has been presented for a subset of PHANGS targets (Sun et al. 2020a; hereafter S20). This Letter provides updated measurements across the full PHANGS–ALMA sample, leveraging the latest processing of the same observational data sets used in Utomo et al. (2018), S20, and Querejeta et al. (2021), as well as applying refined methodologies for converting observable quantities into physical quantities.

2. Data

We base our analysis on the PHANGS high-level measurement database described in Sun et al. (2022, hereafter S22). The database incorporates multiwavelength data for 80 galaxies, extracts observational measurements and associated

Table 1
Methodological Choices

Method Shorthand	α_{CO}	SFR Calibration
“Fiducial”	S20	H α +22 μm
“FUV+W4 SFR”	S20	FUV+22 μm
“Av-corr H α SFR”	S20	A ν -corrected H α ^a
“MW α_{CO} ”	Galactic	H α +22 μm
“B13 α_{CO} ”	B13	H α +22 μm
“G20 α_{CO} ”	G20	H α +22 μm

Note. Reference for all SFR calibrations used here: Belfiore et al. (2023).
a Available only for 19 galaxies in the PHANGS-MUSE sample (Emsellem et al. 2022).

uncertainties with matched sampling and weighting schemes, and converts them into physical quantities following a set of best practices. In this paper, we use the latest version of this database, which sees many improvements over the version published in S22. We summarize these improvements and announce the online release of the associated data products in Appendix A.

Below we list the key physical quantities used in this paper and describe their data sources as well as the methodological choices and assumptions involved in their derivation.

1. *SFR surface density*, Σ_{SFR} . We derive this quantity at a fixed 1.5 kpc resolution⁴¹ from three different SFR calibrations that combine UV, optical, and/or IR data (see Table 1). Our fiducial choice is to combine narrowband H α data (acquired with the 2.5 m du Pont telescope and the ESO/MPG 2.2 m telescope; A. Razza et al. 2023, in preparation) with Wide-field Infrared Survey Explorer (WISE) 22 μm data (Leroy et al. 2019). For this combination, we use a new SFR calibration proposed by Belfiore et al. (2023) that can better mitigate contamination from IR cirrus in the 22 μm band than the classic Calzetti et al. (2007) calibration. Alternatively, we combine Galaxy Evolution Explorer (GALEX) far-UV data (154 nm; Leroy et al. 2019) with WISE 22 μm data, following another new calibration proposed by Belfiore et al. (2023) for similar purposes. Finally, for 19 galaxies in the PHANGS-MUSE sample (Emsellem et al. 2022), we also include H α -based SFR measurements that are corrected for dust extinction based on the Balmer decrement (Belfiore et al. 2023, also see Pessa et al. 2021). This last calibration is likely the most reliable (because of the trustworthy [N II] subtraction, extinction correction, and superior depth of the MUSE observations), but unfortunately, the required MUSE data are available for only $\sim 1/4$ of the galaxies studied here. We note that our H α -based SFR calculations include all diffuse ionized gas emission, as recommended for kiloparsec-scale observations of star-forming galaxies (Belfiore et al. 2022). All our Σ_{SFR} measurements have been corrected for galaxy inclination (as described in S22).
2. *Molecular gas surface density*, Σ_{mol} . We derive this quantity from PHANGS-ALMA CO(2–1) data (Leroy et al. 2021a, 2021b) at 1.5 kpc resolution. More specifically, we use CO(2–1) integrated intensities from

the high-completeness, “broad” moment-0 maps (see Leroy et al. 2021b, for more details) and convert them into molecular gas surface densities using four different prescriptions for the CO-to-H $_2$ conversion factor (α_{CO} , also see Table 1). Our fiducial choice is to combine a varying, metallicity-dependent $\alpha_{\text{CO}(1-0)}$ with a fixed CO line ratio of $R_{21} = 0.65$ (Leroy et al. 2022, but see Section 3.5 for some caveats). Here, the metallicity is inferred from empirical scaling relations and ranges 0.6–1.2 solar for most regions studied in this work (see S20). We also employ three alternative α_{CO} treatments: (1) a fixed, Galactic $\alpha_{\text{CO}(1-0)}$ value of $4.35 M_{\odot} \text{pc}^{-2} (\text{K km s}^{-1})^{-1}$; (2) an empirical $\alpha_{\text{CO}(1-0)}$ calibration depending on metallicity and total (gas + star) mass surface density⁴² (Bolatto et al. 2013, hereafter B13); and (3) a simulation-motivated $\alpha_{\text{CO}(2-1)}$ prescription relying on metallicity, CO integrated intensity, and resolution of the CO measurements⁴³ (Gong et al. 2020, hereafter G20). The last two options are likely the most realistic given the additional physics they intend to capture (e.g., gas excitation; see Section 3.5). Details on the implementation of these α_{CO} prescriptions can be found in S22. The Σ_{mol} values are also corrected for galaxy inclination.

3. *Galactic orbital time*, t_{orb} . We derive this quantity from the rotation curve models based on CO(2–1) kinematics presented in Lang et al. (2020), which exist for 62 of our 80 galaxies. As discussed in S22, we use a set of parametrized model fits to the measured rotation curves (J. Nofech et al. 2023, in preparation), which effectively suppress the fluctuation of rotational velocities across radial bins (due to, e.g., noncircular motions). The orbital time is then determined from the galactocentric radius and the local circular velocity given by the rotation curve models.
4. *Population-averaged molecular cloud free fall time*, \bar{t}_{ff} . We derive this quantity for each 1.5 kpc region by calculating the mass-weighted harmonic mean of the free fall time of all molecular clouds located in that region, as described in S22. In this work, we use the molecular cloud free fall time measured from the 150 pc scale CO(2–1) maps, denoted as $\bar{t}_{\text{ff},150\text{pc}}$. This measurement is available for all 80 galaxies. It adopts a simplifying assumption that the emission in each 150 pc beam originates from a beam-filling, spherical cloud (consistent with Sun et al. 2018, 2020b). We adopt the same conversion factor for \bar{t}_{ff} as for Σ_{mol} .
5. *ISM dynamical equilibrium pressure*, P_{DE} . We derive this quantity on a 1.5 kpc scale by combining the total gas surface density $\Sigma_{\text{gas}} = \Sigma_{\text{mol}} + \Sigma_{\text{atom}}$, stellar mass volume density at the disk midplane ρ_{\star} , and vertical gas velocity dispersion $\sigma_{\text{gas},z}$, closely following S20 (also see Ostriker et al. 2010; Ostriker & Kim 2022):

$$P_{\text{DE}} = \frac{\pi G}{2} \Sigma_{\text{gas}}^2 + \Sigma_{\text{gas}} \sqrt{2G\rho_{\star}} \sigma_{\text{gas},z}. \quad (1)$$

⁴¹ This is the best common resolution achievable for all galaxies in our sample since we rely on WISE 22 μm data.

⁴² While the original B13 prescription also includes a molecular cloud surface density term and it is implemented as such in S20 and S22, we have noticed that it often leads to unphysical α_{CO} values in low surface density regimes. Here we use a fixed cloud surface density of $100 M_{\odot} \text{pc}^{-2}$ to mitigate this issue.

⁴³ We first calculate the G20 α_{CO} based on 150 pc resolution CO data and then spatially average it to 1.5 kpc resolution (see Appendix B in S22).

Here we calculate Σ_{atom} from HI 21 cm line data gathered by various observing programs on the Very Large Array (VLA), Australia Telescope Compact Array (ATCA), and Westerbork Synthesis Radio Telescope (WSRT; see S22 for a full list). We estimate ρ_* by first calculating the 2D stellar mass surface density Σ_* from near-IR data gathered by WISE and Spitzer Space Telescope (with a locally determined mass-to-light ratio; see Leroy et al. 2021a) and then converting it to stellar volume density assuming the stellar disk thickness scales with its radial extent (see Kregel et al. 2002, S20). We adopt a fixed $\sigma_{\text{gas},z} = 11 \text{ km s}^{-1}$ following Ostriker et al. (2010, but see S20 and Ostriker & Kim 2022 for discussions about the caveats related to this assumption). In total, we are able to measure P_{DE} in 48 out of 80 galaxies, with the sample size limited primarily by the availability of HI data.

3. Results

With all the key physical quantities listed in Section 2, we examine the four SF laws described in Section 1 across the full PHANGS–ALMA sample. With over 2000 kiloparsec-size regions across 80 galaxies, this is the largest sample for which all the necessary quantities (including the orbital time, cloud-scale free fall time, and ISM dynamical equilibrium pressure) can be measured directly from observations. Our uniform methodological treatments across the full sample allow for rigorous comparisons between the four SF laws, as well as systematic explorations of how particular methodological choices (Table 1) influence the quantitative results.

3.1. Molecular Kennicutt–Schmidt Relation

We first examine the relationship between surface densities of molecular gas mass and SFR, commonly known as the molecular Kennicutt–Schmidt (mKS) relation. While the original Kennicutt–Schmidt (KS) relation uses the total gas surface density as the independent variable (Kennicutt 1989), it has been shown that the relation with Σ_{mol} is tighter and has a more consistent slope across diverse environments (e.g., Wong & Blitz 2002; Bigiel et al. 2011). This slope is often found to be close to unity in local star-forming galaxies. As a result, the molecular gas depletion time $t_{\text{dep}} \equiv \Sigma_{\text{mol}}/\Sigma_{\text{SFR}}$ varies only weakly, with typical values of 1–3 Gyr (e.g., Leroy et al. 2008; Saintonge et al. 2011; also see review by Saintonge & Catinella 2022).

Figure 1 shows the mKS relation measured at the 1.5 kpc scale (left panel) and the corresponding distribution of t_{dep} (right panel) across the full PHANGS–ALMA sample. With the fiducial methodological choices (see Table 1), our measurements span three decades in Σ_{mol} (10^{-1} – $10^2 M_{\odot} \text{ pc}^{-2}$) and Σ_{SFR} (10^{-4} – $10^{-1} M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$). The corresponding t_{dep} distribution shows a median value and a $\pm 1\sigma$ range of $2.0_{-1.0}^{+1.7}$ Gyr. A large fraction of our measurements are 3σ upper limits⁴⁴ at $\Sigma_{\text{mol}} < 1 M_{\odot} \text{ pc}^{-2}$ or $\Sigma_{\text{SFR}} < 10^{-3} M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$ due to the finite sensitivities of the CO, H α , or IR observations. The distributions of detections and 3σ upper limits in the Σ_{mol} – Σ_{SFR} space also vary moderately depending on the choice of SFR calibrations and α_{CO} prescriptions.

To further quantify the shape and tightness of the mKS relation, we fit a power-law model to the data distribution with a functional form of

$$\log_{10} \left(\frac{\Sigma_{\text{SFR}}}{M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}} \right) = \alpha + \beta \log_{10} \left(\frac{\Sigma_{\text{mol}}}{10 M_{\odot} \text{ pc}^{-2}} \right), \quad (2)$$

where the normalization α is determined at $\Sigma_{\text{mol}} = 10 M_{\odot} \text{ pc}^{-2}$ (close to the midpoint of our sample). The model fit is performed in logarithmic space with the `linmix` package (Kelly 2007). It determines the power-law normalization (α), slope (β), and intrinsic scatter (σ) around it from the data distribution, accounting for measurement uncertainties and nondetections for the dependent variable. We further restrict the fit to measurements above a Σ_{mol} threshold in order to minimize biases caused by nondetections for the independent variable. This threshold is $\Sigma_{\text{mol}} = 1 M_{\odot} \text{ pc}^{-2}$ for the fiducial α_{CO} and varies between 1 and $2 M_{\odot} \text{ pc}^{-2}$ for different α_{CO} choices (see Appendix B).

The first part of Table 2 reports the best-fit model parameters for the mKS relation when adopting various SFR calibrations and α_{CO} . In all cases, we see near-unity power-law indices ($\beta = 0.88$ – 1.21) and small intrinsic scatters ($\sigma = 0.29$ – 0.38 dex). The near-unity slope means the scatter in t_{dep} (about a factor of 2) is almost identical to the residual scatter around the best-fit power-law relation. Considering the impact of SFR calibration and α_{CO} choices separately, we find the former can change the slope by 13% and the normalization by 0.17 dex, whereas the latter produces changes of 25% for the slope and 0.18 dex for the normalization. These findings call special attention to the methodology-dependent nature of the mKS relation shape and slope.

3.2. Molecular Elmegreen–Silk Relation

The KS relation links the current SFR to the amount of gas available at the moment. If there is a characteristic timescale on which the gas is converted into stars, then a potentially more direct (and more physical) relation would connect the current SFR to the amount of gas *normalized by that timescale*. The mES relation (Elmegreen 1997; Silk 1997) is one such candidate, for which the normalizing timescale is the galactic orbital time, t_{orb} . This timescale is relevant to not only galactic rotation but also to shear, spiral arm passages, and cloud–cloud collisions, all of which can regulate cloud formation/destruction and in this way denote the relevant timescale for SF (e.g., Tan 2000). In this case, the independent variable becomes $\Sigma_{\text{mol}}/t_{\text{orb}}$, and the ratio between Σ_{SFR} and this new independent variable defines the SF efficiency per orbital time, $\epsilon_{\text{orb}} = (\Sigma_{\text{SFR}}/\Sigma_{\text{mol}})t_{\text{orb}}$.

Figure 2 shows the mES relation (left panel) and the distribution of ϵ_{orb} (right panel). These measurements are available for a subsample of 62 galaxies, for which we can determine t_{orb} from rotation curve models (see Section 2). In addition, it is often challenging to measure the rotation curve near the edge of the CO images, due to incomplete azimuthal coverage and sparse CO detection. Consequently, there are

⁴⁴ Upper limits of Σ_{mol} are omitted in all figures for clarity.

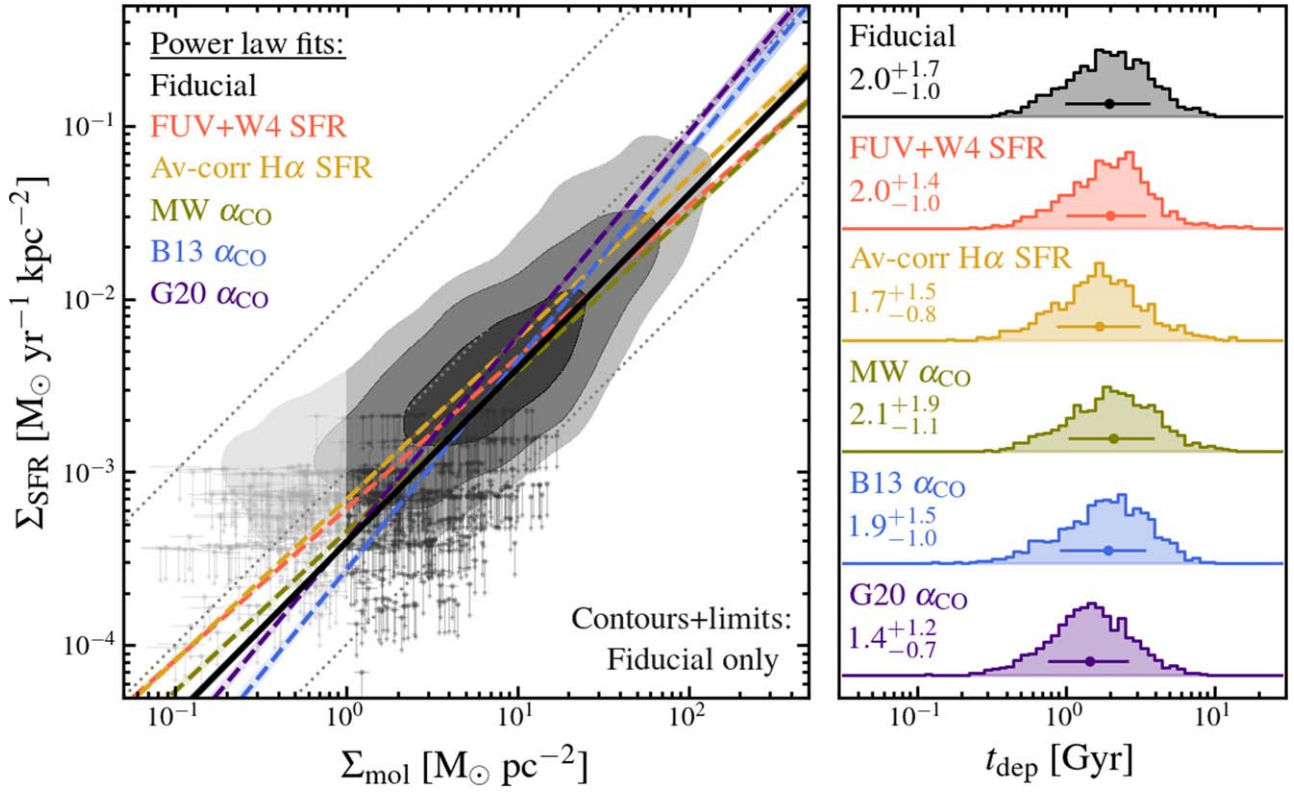


Figure 1. Left panel: the molecular Kennicutt–Schmidt (mKS) relation across the PHANGS–ALMA sample. The density contours (40%–80%–95% levels) show the distributions of all 1.5 kpc scale regions with $>3\sigma$ detections for both Σ_{mol} and Σ_{SFR} , and the downward arrows show 3σ upper limits for Σ_{SFR} (see Section 3). The solid black line shows the best-fit power-law model for all detections and upper limits above $\Sigma_{\text{mol}} = 1 M_{\odot} \text{pc}^{-2}$ (i.e., the nonshaded side for the contours and symbols), where there is minimal censoring on Σ_{mol} . The other colored lines show the best-fit model when using alternative SFR calibrations or CO-to- H_2 conversion factors. The thin dotted lines mark linear relations with constant molecular gas depletion times of 0.1, 1, and 10 Gyr (top left to bottom right). Right panel: normalized histograms of the molecular gas depletion time $t_{\text{dep}} = \Sigma_{\text{mol}}/\Sigma_{\text{SFR}}$, color-coded similarly to the left panel. The median value and 16th–84th percentile range are marked by a dot and a horizontal bar beneath each histogram, with their values displayed to the left of each histogram.

visibly fewer measurements at the low Σ_{SFR} end in Figure 2 than in Figure 1.

Within the range of environments where we do have an adequate number of measurements (i.e., $\Sigma_{\text{SFR}} \sim 10^{-3.5} - 10^{-1} M_{\odot} \text{yr}^{-1} \text{kpc}^{-2}$), we fit a power-law model for the mES relation:

$$\begin{aligned} & \log_{10} \left(\frac{\Sigma_{\text{SFR}}}{M_{\odot} \text{yr}^{-1} \text{kpc}^{-2}} \right) \\ &= \alpha + \beta \log_{10} \left(\frac{\Sigma_{\text{mol}}/t_{\text{orb}}}{0.1 M_{\odot} \text{yr}^{-1} \text{kpc}^{-2}} \right). \end{aligned} \quad (3)$$

The mES relation has a much shallower slope than the mKS relation for any given methodological choice (Table 2). This is expected, as regions with higher Σ_{mol} are typically located at smaller galactocentric radii and thus have shorter t_{orb} . In other words, the dynamic range in $\Sigma_{\text{mol}}/t_{\text{orb}}$ is usually wider than that in Σ_{mol} for the same set of regions, resulting in a shallower mES relation than the mKS relation. While this is not obvious from Figures 1 and 2 due to the different samples of measurements they include, we have verified it in a common subsample (i.e., the intersection of the samples in Figures 1 and 2).

The intrinsic scatter of the mES relation is also smaller than that of the mKS relation for any given methodological choice. This suggests that the mES relation makes empirically better

predictions for Σ_{SFR} than the mKS relation. However, the sublinear slope of the mES relation indicates that ϵ_{orb} decreases systematically toward the high Σ_{SFR} end and shows a wider distribution than t_{dep} across the whole sample. It is thus not a good assumption to adopt a fixed $\epsilon_{\text{orb}} = 5\% - 10\%$ across all regions, even though it could be reasonable to assume a fixed t_{dep} for the same range of conditions.

3.3. Free Fall Time-regulated SF Relation

Aside from the galactic orbital time, another highly relevant timescale for SF is the free fall time of molecular clouds, especially since molecular clouds are the immediate sites of SF. The FFTR relation builds on this notion and connects the local SFR to the ratio of Σ_{mol} and the average free fall time of molecular clouds, \bar{t}_{ff} , in the same region (e.g., Krumholz et al. 2009, 2012). The SF efficiency per free fall time, $\epsilon_{\text{ff}} = (\Sigma_{\text{SFR}}/\Sigma_{\text{mol}})\bar{t}_{\text{ff}}$, then describes the fraction of gas mass converted to stars over a unity \bar{t}_{ff} . This parameter is of particular interest to both observers and theorists, as it can be determined from observable quantities (e.g., Utomo et al. 2018; Evans et al. 2022; see Krumholz et al. 2019 for a compilation) and predicted from analytical and numerical models of turbulence-regulated SF at cloud scales (e.g., Krumholz & McKee 2005; Hennebelle & Chabrier 2011; Federrath & Klessen 2012; Padoan et al. 2012, 2014; Kim et al. 2021).

Figure 3 shows the FFTR relation (left panel) and the distribution of ϵ_{ff} (right panel) across our full sample of 80 galaxies. Our measurements span a similar range in Σ_{SFR} here

Table 2
Best-fit Parameters for the SF Laws

Method	α	β	σ	$N_{\text{det}}/N_{\text{uplim}}$
Molecular Kennicutt–Schmidt Relation (mKS; Section 3.1)				
Fiducial	−2.40	1.00	0.36	1585/450
FUV+W4 SFR	−2.34	0.88	0.29	2279/158
Av-corr H α SFR	−2.23	0.93	0.29	974/0
MW α_{CO}	−2.43	0.92	0.37	1553/451
B13 α_{CO}	−2.36	1.21	0.38	1016/225
G20 α_{CO}	−2.22	1.18	0.35	1298/225
Molecular Elmegreen–Silk Relation (mES; Section 3.2)				
Fiducial	−2.23	0.77	0.31	1001/169
FUV+W4 SFR	−2.19	0.67	0.26	1198/31
Av-corr H α SFR	−2.06	0.78	0.28	516/0
MW α_{CO}	−2.26	0.69	0.33	999/167
B13 α_{CO}	−2.17	0.90	0.32	666/97
G20 α_{CO}	−2.11	0.77	0.33	908/121
Free-fall Time-regulated SF Relation (FFTR; Section 3.3)				
Fiducial	−2.32	0.65	0.34	1457/311
FUV+W4 SFR	−2.28	0.57	0.28	2036/66
Av-corr H α SFR	−2.16	0.62	0.28	880/0
MW α_{CO}	−2.34	0.62	0.34	1439/314
B13 α_{CO}	−2.29	0.75	0.36	970/178
G20 α_{CO}	−2.20	0.76	0.33	1294/220
Pressure-regulated SF Relation (PR; Section 3.4)				
Fiducial	−2.95	0.93	0.33	1138/313
FUV+W4 SFR	−2.84	0.84	0.24	1696/133
Av-corr H α SFR	−2.72	0.85	0.25	651/0
MW α_{CO}	−2.94	0.86	0.33	1122/309
B13 α_{CO}	−2.95	1.08	0.32	1015/224
G20 α_{CO}	−2.87	1.05	0.31	952/171

Note. The values reported here are power-law normalization (α , in dex units), slope (β), intrinsic scatter (σ , in dex units), and number of detections and upper limits used in the model fit ($N_{\text{det}}/N_{\text{uplim}}$). See Equations (2)–(5) for the exact parameterization for each SF law. The formal statistical uncertainties on the best-fit parameters are not listed here as they are negligible compared to the systematic uncertainties associated with SFR calibrations and α_{CO} prescriptions.

as in Figure 1. However, regions with low Σ_{mol} often do not have \bar{t}_{ff} measurements, as they require (a) detecting individual molecular clouds in CO and (b) having enough clouds in the 1.5 kpc region to determine a population average. The range of environments we can probe for the FFTR relation ends up being similar to those used in the power-law fit for the mKS relation (i.e., those above the Σ_{mol} threshold; see Section 3.1).

For the FFTR relation, we fit a power-law model as

$$\begin{aligned} \log_{10} \left(\frac{\Sigma_{\text{SFR}}}{M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}} \right) \\ = \alpha + \beta \log_{10} \left(\frac{\Sigma_{\text{mol}} / \bar{t}_{\text{ff}}}{M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}} \right). \end{aligned} \quad (4)$$

Similar to our findings for the mES relation, the slope of the FFTR relation is also consistently shallower than the slope of the mKS relation, regardless of methodological choices

(Table 2). This reflects a systematic trend that the molecular cloud populations in high Σ_{mol} regions have on average higher densities (e.g., S22) and consequently shorter free fall times. The slope we find for the FFTR relation is sublinear in most cases, implying that ϵ_{ff} drops substantially in higher surface density environments.

The FFTR relation exhibits an intrinsic scatter of $\sigma \approx 0.3$ dex, which is comparable to the mES relation and mildly smaller than the mKS relation. But the 1σ dispersion of the ϵ_{ff} distribution is wider, again due to the sublinear FFTR relation slope. That is, the measured range of $\epsilon_{\text{ff}} = 0.7^{+1.1}_{-0.4}\%$ (for the fiducial SFR calibration and α_{CO}) can be seen as the combined results of the FFTR relation intrinsic scatter *plus* a systematic trend of decreasing ϵ_{ff} with Σ_{SFR} (and Σ_{mol}). For studies that rely on assumed constant ϵ_{ff} values to predict SFR (as is done in many galaxy simulations), it would be important to also account for this systematic trend.

In the context of turbulence-regulated SF models, variations in ϵ_{ff} are considered to be driven by changes in the physical properties of individual star-forming molecular clouds (e.g., Krumholz & McKee 2005; Federrath & Klessen 2012). Broadly speaking, ϵ_{ff} would be higher for clouds with higher turbulent Mach number \mathcal{M} (which is proportional to the turbulent velocity dispersion σ_{turb}) and lower virial parameter α_{vir} (e.g., see Figure 1 in Federrath & Klessen 2012). Since high Σ_{mol} and Σ_{SFR} regions tend to host molecular clouds with larger σ_{turb} and smaller α_{vir} (see Figure B1 in S22), we would then expect ϵ_{ff} to be higher in those regions. Yet this expectation appears inconsistent with the empirical trends found in this work and in previous studies (e.g., Leroy et al. 2017; Schrubba et al. 2019; but see Barnes et al. 2017 for caveats). That being said, a more rigorous and thorough comparison with theoretical predictions is beyond the scope of this work and will be addressed in S. Meidt et al. 2023 (in preparation).

3.4. Pressure-regulated SF Relation

The mES and FFTR relations discussed above measure the SF efficiency relative to a specific dynamical timescale (either t_{orb} or \bar{t}_{ff}). The focus, explicitly or implicitly, is on the “mass supply” aspect, with SF thought of as a process that depletes the ISM. The pressure-regulated, feedback-modulated SF theory (Ostriker et al. 2010; Ostriker & Kim 2022) instead views SF as a source of energy and momentum, rather than a sink of mass, for the ISM. In this framework, the local SFR determines the energy and momentum injection rate into the ISM via stellar and supernovae feedback, which over time offsets turbulence dissipation and radiative cooling and prevents the ISM from collapsing in the galactic gravitational potential (see also Thompson et al. 2005, who propose that feedback maintains the ISM in a marginally Toomre-stable state). The local SFR required to keep the ISM in a long-term equilibrium is thus ultimately set by the weight of the ISM in the galactic potential. Since we expect the ISM in massive, star-forming disk galaxies to exist in such a thermal and dynamical equilibrium, it implies a proportionality between Σ_{SFR} and the weight of the ISM per unit area, commonly referred to as the dynamical equilibrium pressure, P_{DE} . Their ratio $\Upsilon_{\text{fb}} = P_{\text{DE}}/\Sigma_{\text{SFR}}$ is named the feedback yield, as it quantifies the ISM pressure resulting from the injection of momentum and energy by feedback, measured per unit stellar mass formed (see Ostriker & Kim 2022, for a recent formulation of this theory).

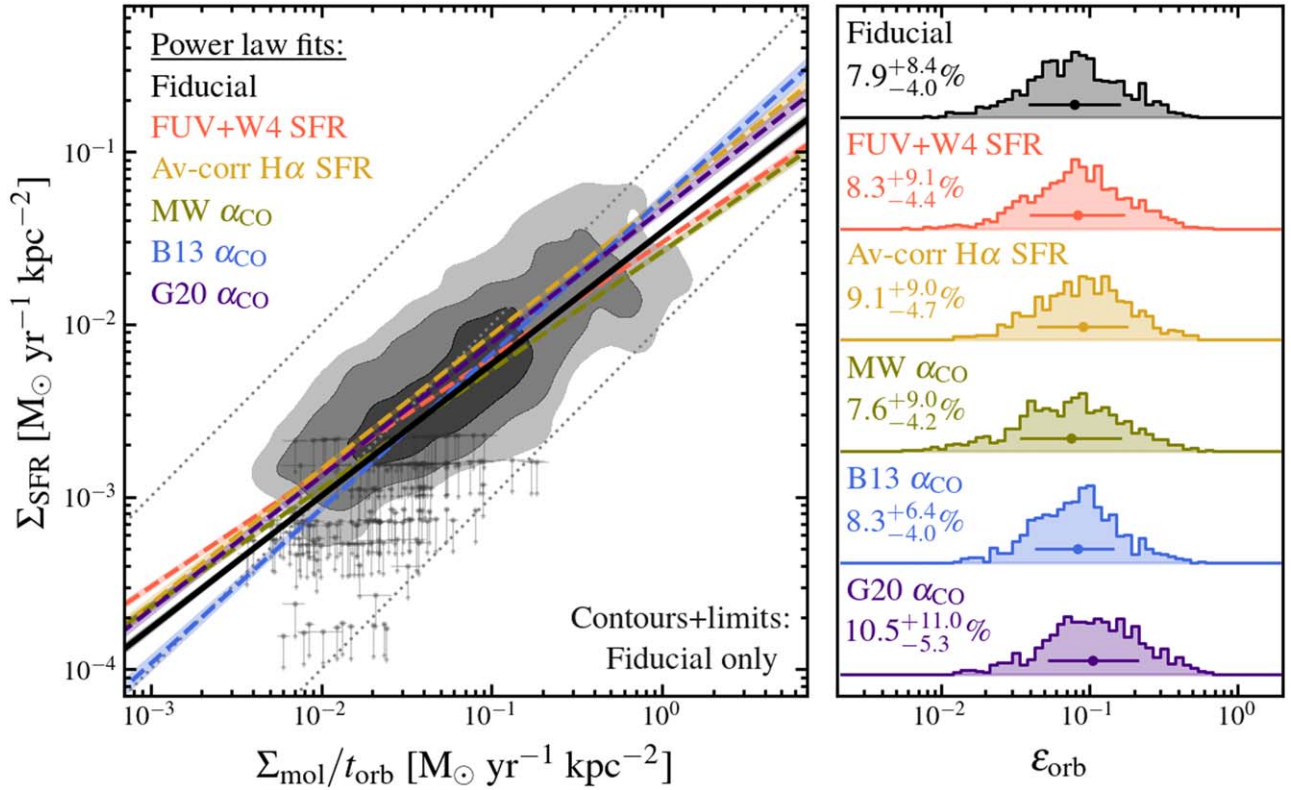


Figure 2. Similar to Figure 1, but here showing the mES relation (left panel) and normalized histograms of the SF efficiency per orbital time, $\epsilon_{\text{orb}} = (\Sigma_{\text{SFR}}/\Sigma_{\text{mol}})t_{\text{orb}}$ (right panel). The thin dotted lines in the left panel mark linear relations with constant ϵ_{orb} of 1%, 10%, and 100% (bottom right to top left). Note that data below the Σ_{mol} threshold shown in Figure 1 are also excluded in the power-law fit for the mES (and all other) relations.

Figure 4 shows the $\Sigma_{\text{SFR}}-P_{\text{DE}}$ relation, that is, the PR relation, and the distribution of the Υ_{fb} parameter. This analysis covers a subsample of 48 galaxies that have HI 21 cm data available to us (see Section 2 and S22). The measurements shown here represent a major update over those presented in S20, which only covered 28 galaxies and relied on earlier versions of the PHANGS-ALMA CO data and associated HI data.

We fit a power-law model to the PR relation with the following parameterization:

$$\log_{10} \left(\frac{\Sigma_{\text{SFR}}}{M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}} \right) = \alpha + \beta \log_{10} \left(\frac{P_{\text{DE}}}{10^4 k_{\text{B}} \text{ K cm}^{-3}} \right). \quad (5)$$

Across our sample, the PR relation exhibits a mildly sublinear slope for several methodological choices (including the fiducial), which is broadly consistent with the results reported in S20 and in other studies (e.g., Fisher et al. 2019, 2022; Barrera-Ballesteros et al. 2021). When adopting the B13 or G20 α_{CO} , the slope appears near unity, although still not as steep as that seen in numerical simulations ($\beta = 1.1-1.2$; see Kim et al. 2013; Ostriker & Kim 2022). This small discrepancy may be due to a limited range of conditions ($\Sigma_{\text{SFR}} \approx 10^{-3}-10^{-1} M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$) probed in the observations. Alternatively, it could be attributed to (1) an elevated Υ_{fb} in high Σ_{SFR} regions, possibly caused by more efficient feedback from clustered supernovae in reality (e.g.,

Fisher et al. 2019, 2022; Gentry et al. 2019), or (2) additional sources of turbulence injection, such as gas radial inflows (e.g., Krumholz et al. 2018; Girard et al. 2021). Otherwise, the observed Υ_{fb} range of $1-3 \times 10^3 \text{ km s}^{-1}$ is broadly in line with simulation results (e.g., Kim et al. 2013, 2017; Keller et al. 2022; Ostriker & Kim 2022).

The intrinsic scatter around the PR relation is mildly smaller than that around the mKS relation and more comparable to the other two relations (Table 2). This means the PR relation makes as good empirical predictions for Σ_{SFR} as the mES and FFTR relations, but with similar limitations given its possibly sublinear slope and thus a varying Υ_{fb} . Practical applications of the PR relation for this purpose should also consider the systematic trends in Υ_{fb} with Σ_{SFR} and other environmental conditions.

We note that among the four SF laws examined in this work, only the PR relation includes both the molecular and the atomic gas. The first three relations focus only on the molecular gas, which means that they implicitly take the atomic-to-molecular phase balance as given and do not capture any physics related to that process. In contrast, the PR relation *has* to include both phases because it concerns the energy and momentum budget, which has no natural border between the phases. To some degree, this makes the PR relation more generally applicable, even to regions or galaxies with no detectable molecular gas (e.g., Kado-Fong et al. 2022). Nevertheless, the need for HI data limits our ability to study the PR relation for a larger fraction of the PHANGS-ALMA sample. This situation will likely improve as we gather more HI data for PHANGS targets with VLA and MeerKAT (A. Sardone et al. in preparation; C. Eibensteiner et al. in preparation).

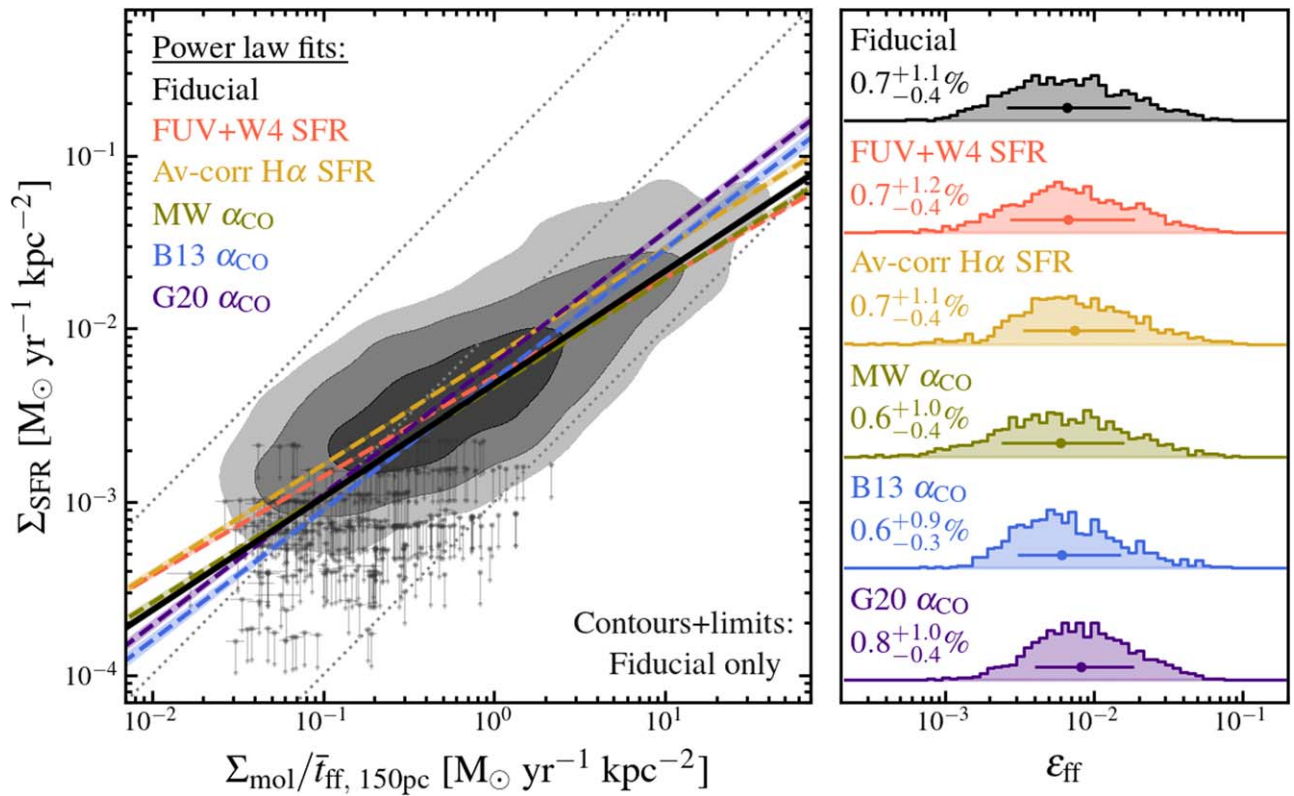


Figure 3. Similar to Figure 1, but here showing the FFTR relation (left panel) and normalized histograms of the SF efficiency per free fall time, $\epsilon_{\text{ff}} = (\Sigma_{\text{SFR}}/\Sigma_{\text{mol}})\bar{\tau}_{\text{ff}}$ (right panel). The thin dotted lines in the left panel mark linear relations with constant ϵ_{ff} of 0.1%, 1%, and 10% (bottom right to top left).

3.5. Impacts of Methodological Choices

For all four SF laws, we observe coherent, systematic changes in their shapes when we adopt different methodological choices (Table 2). Changes in the best-fit power-law parameters due to methodological choices are much larger than their formal statistical uncertainties, which suggests that methodology-related systematics are clearly a dominant source of uncertainties. Here we briefly summarize these systematic changes and discuss their implications. Note that while the number of available measurements differs for each choice (Table 2), the trends we see remain the same when doing rigorous comparisons with a matched sample of measurements.⁴⁵

Considering the impacts of both SFR calibrations and α_{CO} prescriptions, we see that the former can introduce at least a 10%–15% change in the slopes and a 0.15–0.25 dex variation in the normalization factors of the SF laws.⁴⁶ The latter can produce a 20%–25% difference in the slopes and 0.10–0.20 dex in the normalization. All these systematics reflect real uncertainties in our empirical knowledge of ISM physics (such as heating and shielding) and its interactions with stellar populations.

Between the three SFR calibrations, we find that the FUV +22 μm calibration yields slightly higher Σ_{SFR} values at the low end compared to the fiducial H α +22 μm calibration, resulting in shallower SF law slopes. This is in line with the

findings by Belfiore et al. (2023) that the former is still somewhat more susceptible to contamination from IR cirrus at $\Sigma_{\text{SFR}} \lesssim 10^{-3} M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$, even though both calibrations have seen substantial improvements in this aspect compared with previous versions (Leroy et al. 2019). The MUSE extinction-corrected H α calibration gives more consistent SF law slopes with the fiducial H α +22 μm , but it yields a slightly higher SF law normalization (by 0.1–0.2 dex). This discrepancy can be traced back to an intrinsic zero-point difference between H α maps obtained from MUSE (Emsellem et al. 2022) versus narrowband observations (A. Razza et al. in preparation), an issue to be addressed in future processing of the narrowband data.

Between the four α_{CO} prescriptions, we find that the fiducial S20 α_{CO} yields higher Σ_{mol} at the low end than the Galactic α_{CO} , as expected from its built-in metallicity dependence. The B13 prescription produces even higher α_{CO} at the low end due to its stronger, exponential metallicity dependence; it also produces low α_{CO} in high surface density regions, which comes from an extra negative dependence on the total (stellar + gas) surface density that aims to account for elevated gas excitation and velocity dispersion in the “starburst” regime. G20 gives similarly low α_{CO} at the high surface density end as B13, which is expected because the extra dependence on CO intensity in the G20 α_{CO} aims to capture a similar set of physics. At the low surface density end, the G20 α_{CO} is closer to the S20 α_{CO} as it also features a power-law metallicity dependence. Overall, we expect the B13 and G20 prescriptions to likely yield more realistic α_{CO} than the other two options given the extra physics they (at least intend to) capture. That being said, the exact behavior of α_{CO} in the high-density, high-excitation, “starburst” regime and the functional

⁴⁵ We do not show the quantitative results here for the sake of brevity but note that these results can be easily reproduced with the published data products described in Appendix A.

⁴⁶ These likely represent only lower limits because the three SFR prescriptions considered here were calibrated in the same way (Belfiore et al. 2023) and thus have some built-in uniformity.

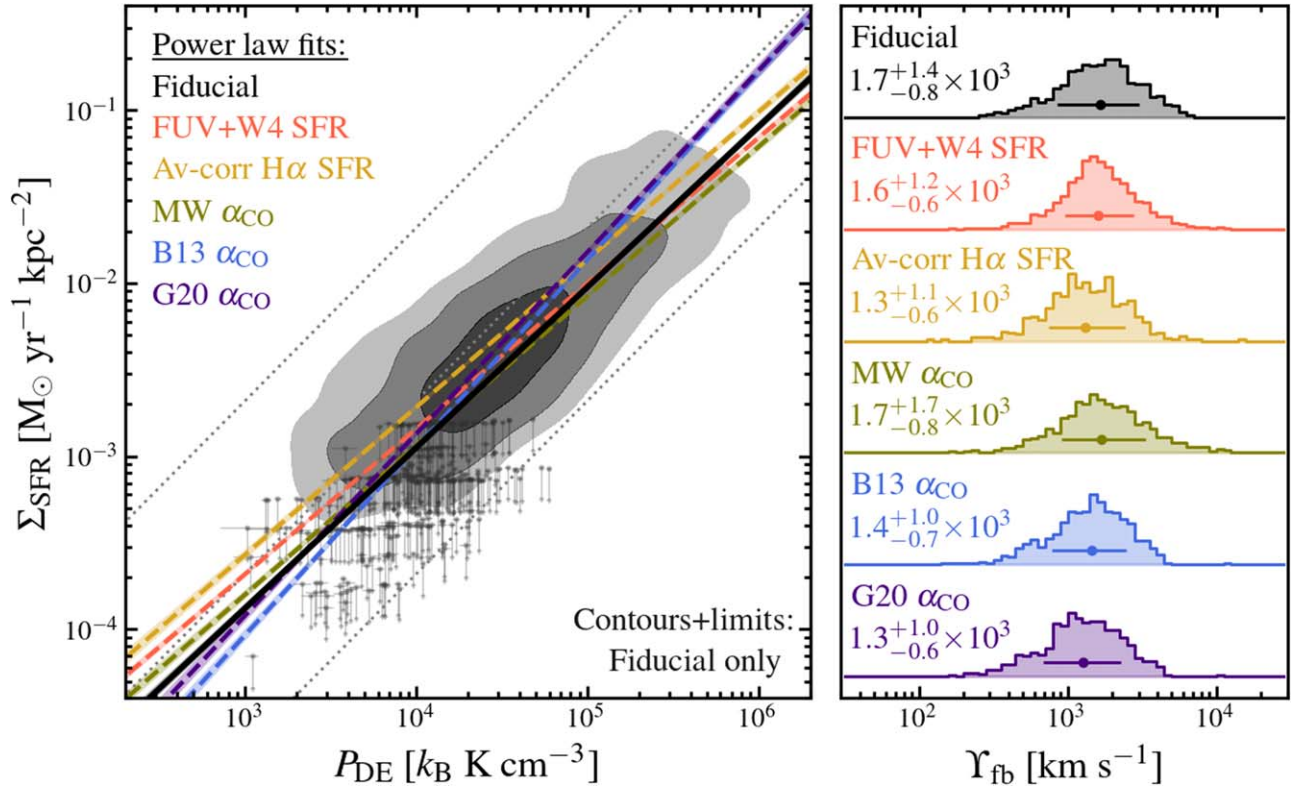


Figure 4. Similar to Figure 1, but here showing the PR relation (left panel) and normalized histograms of the feedback yield, $\Upsilon_{\text{fb}} = P_{\text{DE}}/\Sigma_{\text{SFR}}$ (right panel). The thin dotted lines in the left panel mark linear relations with constant Υ_{fb} of 10^2 , 10^3 , and 10^4 km s^{-1} (top left to bottom right).

form of its metallicity dependence both remain key open questions.

Related to our treatment of α_{CO} , our fiducial treatment uses a constant CO 2–1/1–0 ratio of $R_{21} = 0.65$ in combination with the Galactic, S20, and B13 $\alpha_{\text{CO}(1-0)}$ values.⁴⁷ The assumption of a constant R_{21} is not realistic on its own, especially given recent findings of an approximate scaling of $R_{21} \propto \Sigma_{\text{SFR}}^{0.15}$ by several studies (den Brok et al. 2021; Yajima et al. 2021; Leroy et al. 2022, 2023). If we combine such a Σ_{SFR} -dependent R_{21} value with the fiducial $\alpha_{\text{CO}(1-0)}$, the SF law slope would increase from $\beta = 1.00$ to $\beta/(1 - 0.15\beta) \approx 1.18$, thus agreeing better with the results found when using the B13 or G20 α_{CO} . While further investigation on R_{21} is beyond the scope of this work, we expect to improve our fiducial α_{CO} treatment in the near future by explicitly incorporating the R_{21} prescription suggested in Leroy et al. (2022, 2023).

4. Conclusions

In this Letter, we examine four SF laws commonly used in the literature based on resolved, multiwavelength measurements of SFR and ISM properties across 80 nearby galaxies. This work represents a major improvement over several previous studies (e.g., Wong & Blitz 2002; Bigiel et al. 2011; Leroy et al. 2013; Utomo et al. 2018; Sun et al. 2020a; Querejeta et al. 2021) given the larger number of galaxies covered, the higher quality of the underlying observational data, and the consistent methodological treatments for deriving

relevant physical quantities (especially SFR and molecular gas mass).

We measure the slopes, normalization factors, and intrinsic scatters of the four SF laws as well as their dependence on methodological choices. We also report the distributions of the “proportionality constants” for the SF laws, which quantify SF efficiencies and feedback yield. Our key findings are as follows:

1. Within the range of conditions probed in our sample, each of the four SF laws is well described by a single power law, with typical intrinsic scatter of 0.3–0.4 dex. For any given set of methodological choices, the mKS relation consistently shows a $\sim 10\%$ larger intrinsic scatter than the other three SF laws. The latter relations can thus provide slightly better empirical predictions for the local SFR surface density.
2. Modulo systematic uncertainties related to methodological choices, we find a near-unity slope ($\beta \approx 0.9\text{--}1.2$) for the mKS relation, which implies a roughly constant molecular gas depletion time of 1–3 Gyr. The mES relation and the FFTR relation both have sublinear slopes ($\beta \approx 0.6\text{--}0.9$) for most methodological choices, which means that the SFE per orbital time (typically 5%–10%) and the SFE per free fall time (typically 0.5%–1%) both become lower under higher surface density conditions. The PR relation is also mildly sublinear in most cases ($\beta \approx 0.8\text{--}1.0$), signifying a potential increase in the feedback yield (typically $1\text{--}3 \times 10^3 \text{ km s}^{-1}$) or possible contributions from other turbulence driving mechanisms in high surface density environments.
3. The exact shapes of the SF laws and the distribution functions of physical parameters vary systematically with

⁴⁷ The B13 α_{CO} was calibrated primarily with CO(2–1) observations assuming a fixed R_{21} , so it should not be combined with a varying R_{21} . The G20 α_{CO} was also explicitly calibrated for the CO(2–1) transition and thus does not need an assumed R_{21} .

the adopted SFR calibration and CO-to-H₂ conversion factor. The former introduces at least a 10%–15% uncertainty on the SF law slopes and 0.15–0.25 dex on the normalization, whereas the latter produces differences of 20%–25% for the slopes and 0.10–0.20 dex for the normalization. This is a general issue applicable to not only the SFR calibrations and conversion factors examined in this work but also to others used in the literature. It remains the dominant source of uncertainties for observational studies of SF laws. Comparisons between different observational data sets as well as theoretical interpretations hinging heavily on the SF law slopes should be particularly cautious about these systematics.

Looking forward, we expect studies that probe more extreme environmental conditions (e.g., very high or very low surface densities) to reveal possible changes of behavior in any of the SF laws, which would indicate additional physics at play. With much observational efforts already devoted to these directions (e.g., Daddi et al. 2010; Genzel et al. 2010; de los Reyes & Kennicutt 2019; Wilson et al. 2019; Kennicutt & De Los Reyes 2021; Fisher et al. 2022; Kado-Fong et al. 2022), the next critical step would be to build a large, homogeneous data set to which one can apply improved, consistent methodological treatments. To make the most of these measurements, it will also be critical to compare the systematic trends that we do observe in the SFE per orbital time, SFE per free fall time, and feedback yield to the results of current analytical and numerical models so as to understand how well these models can describe real star-forming galaxies (e.g., Ostriker & Kim 2022, S.Meidt et al. 2023, in preparation). Finally, with reliable measurements of the molecular disk scale height in PHANGS galaxies (Jefferson et al. 2022; J. Sun et al. 2023, in preparation), we will expand our analysis in this work to cover the volumetric SF laws originally suggested by Schmidt (1959) and actively reconsidered in recent works (e.g., Bacchini et al. 2019).

This work was carried out as part of the PHANGS collaboration. We thank the anonymous referee for constructive feedback and S. Ellison for helpful discussions.

J.S. acknowledges support by the Natural Sciences and Engineering Research Council of Canada (NSERC) through a Canadian Institute for Theoretical Astrophysics (CITA) National Fellowship. The work of A.K.L. is partially supported by the National Science Foundation (NSF) under grant Nos. 1615105, 1615109, and 1653300. The work of E.C.O. was partly supported by grant No. 510940 from the Simons Foundation. E.R. acknowledges the support of NSERC, funding reference number RGPIN-2022-03499. The research of C.D.W. is supported by grants from NSERC and the Canada Research Chairs program. G.A.B. acknowledges the support from ANID Basal project FB210003. A.H. was supported by the Programme National Cosmologie et Galaxies (PNCG) of CNRS/INSU with INP and IN2P3, cofunded by CEA and CNES, and by the Programme National “Physique et Chimie du Milieu Interstellaire” (PCMI) of CNRS/INSU with INC/INP, cofunded by CEA and CNES. E.W.K. acknowledges support from the Smithsonian Institution as a Submillimeter Array (SMA) Fellow and the Natural Sciences and Engineering Research Council of Canada. K.K. gratefully acknowledges funding from the Deutsche Forschungsgemeinschaft (DFG) in

the form of an Emmy Noether Research Group (grant No. KR4598/2-1, PI Kreckel). H.A.P. acknowledges support by the National Science and Technology Council of Taiwan under grant 110-2112-M-032-020-MY3. J.P. acknowledges support by the DAOISM grant ANR-21-CE31-0010 and by the Programme National “PCMI” of CNRS/INSU with INC/INP, cofunded by CEA and CNES. M.Q. acknowledges support from the Spanish grant PID2019-106027GA-C44, funded by MCIN/AEI/10.13039/501100011033. A.S. is supported by an NSF Astronomy and Astrophysics Postdoctoral Fellowship under award AST-1903834. A.U. acknowledges support from the Spanish grants PGC2018-094671-B-I00, funded by MCIN/AEI/10.13039/501100011033 and by “ERDF A Way of Making Europe,” and PID2019-108765GB-I00, funded by MCIN/AEI/10.13039/501100011033. E.S. and T.G.W. acknowledge funding from the European Research Council (ERC) under the European Union’s Horizon 2020 research and innovation program (grant agreement No. 694343). F.Bi acknowledges funding from the ERC under the European Union’s Horizon 2020 research and innovation program (grant agreement No. 726384/Empire). A.D.B. acknowledges partial support from NSF-AST2108140. M.C. gratefully acknowledges funding from the DFG through an Emmy Noether Research Group (grant No. CH2137/1-1). COOL Research DAO is a Decentralized Autonomous Organization supporting research in astrophysics aimed at uncovering our cosmic origins. J.G. gratefully acknowledges financial support from the Swiss National Science Foundation (grant No. CRSII5_193826). S.C.O.G. and R.S.K. acknowledge funding from the ERC via the Synergy Grant “ECOGAL” (project ID 855130). They also acknowledge funding from the DFG via the Collaborative Research Center (SFB 881—138713538), “The Milky Way System” (subprojects A1, B1, B2, and B8), from the Heidelberg Cluster of Excellence (EXC 2181-390900948) “STRUCTURES,” and from the German Ministry for Economic Affairs and Climate Action in project “MAINN” (funding ID 50002206). K.G. is supported by the Australian Research Council through the Discovery Early Career Researcher Award (DECRA) Fellowship DE220100766 funded by the Australian Government. K.G. is supported by the Australian Research Council Centre of Excellence for All Sky Astrophysics in 3 Dimensions (ASTRO 3D), through project No. CE170100013. J.M.D.K. gratefully acknowledges funding from the ERC under the European Union’s Horizon 2020 research and innovation program via the ERC Starting Grant MUSTANG (grant No. 714907).

This paper makes use of the following ALMA data, which have been processed as part of the PHANGS–ALMA CO(2–1) survey:

ADS/JAO.ALMA#2012.1.00650.S, ADS/JAO.ALMA#2013.1.00803.S, ADS/JAO.ALMA#2013.1.01161.S, ADS/JAO.ALMA#2015.1.00121.S, ADS/JAO.ALMA#2015.1.00782.S, ADS/JAO.ALMA#2015.1.00925.S, ADS/JAO.ALMA#2015.1.00956.S, ADS/JAO.ALMA#2016.1.00386.S, ADS/JAO.ALMA#2017.1.00392.S, ADS/JAO.ALMA#2017.1.00766.S, ADS/JAO.ALMA#2017.1.00886.L, ADS/JAO.ALMA#2018.1.01321.S, ADS/JAO.ALMA#2018.1.01651.S, ADS/JAO.ALMA#2018.A.00062.S. ALMA is a partnership of ESO (representing its member states), NSF (USA), and NINS (Japan), together with NRC (Canada), NSC and ASIAA (Taiwan), and KASI (Republic of Korea), in cooperation with the Republic of Chile. The Joint ALMA Observatory is operated by ESO, AUI/

NRAO, and NAOJ. The National Radio Astronomy Observatory (NRAO) is a facility of NSF operated under cooperative agreement by Associated Universities, Inc (AUI).

This work is based in part on observations made with NSF’s Karl G. Jansky Very Large Array (VLA; project code: 14A-468, 14B-396, 16A-275, 17A-073, 18B-184). VLA is also operated by NRAO.

This work is based in part on observations made with the ATCA. ATCA is part of the Australia Telescope National Facility, which is funded by the Australian Government for operation as a National Facility managed by CSIRO.

This work is based in part on observations made with the Westerbork Synthesis Radio Telescope (WSRT) owned by ASTRON. ASTRON, the Netherlands Institute for Radio Astronomy, is an institute of the Dutch Research Council (De Nederlandse Organisatie voor Wetenschappelijk Onderzoek, NWO).

This work is based in part on observations collected at the European Southern Observatory under ESO programs 094.C-0623 (PI: Kreckel), 095.C-0473, 098.C-0484 (PI: Blanc), 1100.B-0651 (PHANGS-MUSE; PI: Schinnerer), 094.B-0321 (MAGNUM; PI: Marconi), 099.B-0242, 0100.B-0116, 098.B-0551 (MAD; PI: Carollo), and 097.B-0640 (TIMER; PI: Gadotti).

This work is based in part on observations made with the Spitzer Space Telescope, which is operated by the Jet Propulsion Laboratory, California Institute of Technology, under a contract with NASA.

This work makes use of data products from the Wide-field Infrared Survey Explorer (WISE), which is a joint project of the University of California, Los Angeles, and the Jet Propulsion Laboratory/California Institute of Technology, funded by NASA.

This work is based in part on observations made with the Galaxy Evolution Explorer (GALEX). GALEX is a NASA Small Explorer, whose mission was developed in cooperation with the Centre National d’Etudes Spatiales (CNES) of France and the Korean Ministry of Science and Technology. GALEX is operated for NASA by the California Institute of Technology under NASA contract NAS5-98034.

This work is based in part on data gathered with the CIS 2.5 m Irénée du Pont Telescope and the ESO/MPG 2.2 m Telescope at Las Campanas Observatory, Chile.

This work has made use of the NASA/IPAC Infrared Science Archive (IRSA) and the NASA/IPAC Extragalactic Database (NED), which are funded by NASA and operated by the California Institute of Technology.

We acknowledge the usage of the SAO/NASA Astrophysics Data System.

Facilities: ALMA, VLA, ATCA, WSRT, VLT:Yepun, Spitzer, WISE, GALEX, Du Pont, Max Planck:2.2 m, IRSA.

Software: NumPy (Harris et al. 2020), Astropy (Astropy Collaboration et al. 2013, 2018), Matplotlib (Hunter 2007), linmix (Kelly 2007), MegaTable (Sun 2022), adstex (<https://github.com/yymao/adstex>).

Appendix A Data Products

This paper uses the latest version of the PHANGS high-level measurement database (S22). This evolving database incorporates homogenized measurements from a rich set of

multiwavelength observations and enables rigorous comparisons across all galaxies. S22 already presented the database construction methodologies (such as observational data sampling and weighing, conversion to physical quantities) in great detail. Here we briefly describe the major improvements we have incorporated since the first publication of the database (v3.0) with S22.

A first major improvement is the ingestion of data products from the PHANGS–MUSE observations (Emsellem et al. 2022). In particular, the analysis in this paper involves the H α line intensity maps and the associated uncertainty maps, which are part of the PHANGS–MUSE Data Analysis Pipeline high-level products (or DAP products; see Section 5 in Emsellem et al. 2022). We correct these maps for dust extinction pixel-by-pixel based on the observed intensity ratio of H α and H β (which is also part of the DAP products). We then calculate the average H α surface brightness over each 1.5 kpc region from both the uncorrected and corrected H α maps. These average surface brightness values, along with their uncertainties (from Gaussian error propagation), are recorded in our high-level database and used to further derive SFR surface densities and their uncertainties (see Section 2).

A second improvement is on the HI 21 cm line data that are used for measuring the atomic gas surface density. The older v3.0 database already incorporated HI data from a variety of surveys (such as VLA:THINGS–Walter et al. 2008, VLA:VIVA–Chung et al. 2009, ATCA:LVHIS–Koribalski et al. 2018, and VLA:PHANGS–A. Sardone et al. in preparation), and all HI measurements were derived from the “official” moment maps produced by each of the survey teams. However, each team created these moment maps from the original data cubes in a slightly different way, which led to an extra layer of inhomogeneity among these data products. We have now reprocessed all HI data starting from their raw data cubes, adopting a uniform set of treatments for cube postprocessing, signal masking, and moment map generation across all surveys. We have also processed additional HI data from the HALOGAS (Heald et al. 2011) and WHISP (van der Hulst et al. 2001) surveys and added them into the mix. These improvements allow us to include HI measurements for more galaxies and allow fairer comparisons between galaxies covered by different surveys.

A third improvement is on the data averaging scheme when extracting measurements for each 1.5 kpc area. For the older v3.0 database, we extracted “area-weighted” average values by directly averaging over all pixels in a native resolution image that fall inside the (sharp) boundary of each 1.5 kpc aperture. This approach was well suited for the analyses in S22, which required maximally independent measurements between adjacent apertures and consistent averaging schemes between cloud-scale molecular gas measurements and large-scale environmental measurements (see Section 3 therein). However, this approach can lead to slightly different “effective smoothing scales” among data sets and galaxies, as the image native resolution varies from case to case even though the averaging aperture size is fixed. Such behavior is not ideal for the analyses in this work because the SF laws are known to vary as a function of spatial scales (e.g., Kreckel et al. 2018; Pessa et al. 2021). Consequently, for all “aperture-scale averaged” measurements used in this work, we instead calculate them by first convolving the native resolution images to a matched Gaussian beam with an FWHM of 1.5 kpc and then directly extracting values from the convolved images at the center of each aperture. This alternative averaging scheme ensures all important quantities used in this work (such as Σ_{SFR} ,

Table A1
Column Descriptions for the Machine-readable Table

Column Name	Unit	Description
gal_name		Galaxy name
RA	°	R.A. of the aperture center
DEC	°	Decl. of the aperture center
Sigma_SFR_HaW4recal	$M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$	SFR surface density ($H\alpha+22 \mu\text{m}$)
e_Sigma_SFR_HaW4recal	$M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$	Error on SFR surface density ($H\alpha+22 \mu\text{m}$)
Sigma_SFR_FUVW4recal	$M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$	SFR surface density (FUV+22 μm)
e_Sigma_SFR_FUVW4recal	$M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$	Error on SFR surface density (FUV+22 μm)
Sigma_SFR_Hacorr	$M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$	SFR surface density (A_V -corrected $H\alpha$)
e_Sigma_SFR_Hacorr	$M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$	Error on SFR surface density (A_V -corrected $H\alpha$)
Sigma_mol_S20	$M_{\odot} \text{ pc}^{-2}$	Molecular gas surface density (S20 α_{CO})
e_Sigma_mol_S20	$M_{\odot} \text{ pc}^{-2}$	Error on molecular gas surface density (S20 α_{CO})
Sigma_mol_MW	$M_{\odot} \text{ pc}^{-2}$	Molecular gas surface density (Galactic α_{CO})
e_Sigma_mol_MW	$M_{\odot} \text{ pc}^{-2}$	Error on molecular gas surface density (Galactic α_{CO})
Sigma_mol_B13	$M_{\odot} \text{ pc}^{-2}$	Molecular gas surface density (B13 α_{CO})
e_Sigma_mol_B13	$M_{\odot} \text{ pc}^{-2}$	Error on molecular gas surface density (B13 α_{CO})
Sigma_mol_G20	$M_{\odot} \text{ pc}^{-2}$	Molecular gas surface density (G20 α_{CO})
e_Sigma_mol_G20	$M_{\odot} \text{ pc}^{-2}$	Error on molecular gas surface density (G20 α_{CO})
Sigma_mol_per_t_orb_S20	$M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$	Molecular gas surface density per orbital time (S20 α_{CO})
e_Sigma_mol_per_t_orb_S20	$M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$	Error on molecular gas surface density per orbital time (S20 α_{CO})
Sigma_mol_per_t_orb_MW	$M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$	Molecular gas surface density per orbital time (Galactic α_{CO})
e_Sigma_mol_per_t_orb_MW	$M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$	Error on molecular gas surface density per orbital time (Galactic α_{CO})
Sigma_mol_per_t_orb_B13	$M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$	Molecular gas surface density per orbital time (B13 α_{CO})
e_Sigma_mol_per_t_orb_B13	$M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$	Error on molecular gas surface density per orbital time (B13 α_{CO})
Sigma_mol_per_t_orb_G20	$M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$	Molecular gas surface density per orbital time (G20 α_{CO})
e_Sigma_mol_per_t_orb_G20	$M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$	Error on molecular gas surface density per orbital time (G20 α_{CO})
Sigma_mol_per_t_ff_S20	$M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$	Molecular gas surface density per free fall time (S20 α_{CO})
e_Sigma_mol_per_t_ff_S20	$M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$	Error on molecular gas surface density per free fall time (S20 α_{CO})
Sigma_mol_per_t_ff_MW	$M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$	Molecular gas surface density per free fall time (Galactic α_{CO})
e_Sigma_mol_per_t_ff_MW	$M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$	Error on molecular gas surface density per free fall time (Galactic α_{CO})
Sigma_mol_per_t_ff_B13	$M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$	Molecular gas surface density per free fall time (B13 α_{CO})
e_Sigma_mol_per_t_ff_B13	$M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$	Error on molecular gas surface density per free fall time (B13 α_{CO})
Sigma_mol_per_t_ff_G20	$M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$	Molecular gas surface density per free fall time (G20 α_{CO})
e_Sigma_mol_per_t_ff_G20	$M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$	Error on molecular gas surface density per free fall time (G20 α_{CO})
P_DE_S20	$k_B \text{ K cm}^{-3}$	ISM dynamical equilibrium pressure (S20 α_{CO})
e_P_DE_S20	$k_B \text{ K cm}^{-3}$	Error on ISM dynamical equilibrium pressure (S20 α_{CO})
P_DE_MW	$k_B \text{ K cm}^{-3}$	ISM dynamical equilibrium pressure (Galactic α_{CO})
e_P_DE_MW	$k_B \text{ K cm}^{-3}$	Error on ISM dynamical equilibrium pressure (Galactic α_{CO})
P_DE_B13	$k_B \text{ K cm}^{-3}$	ISM dynamical equilibrium pressure (B13 α_{CO})
e_P_DE_B13	$k_B \text{ K cm}^{-3}$	Error on ISM dynamical equilibrium pressure (B13 α_{CO})
P_DE_G20	$k_B \text{ K cm}^{-3}$	ISM dynamical equilibrium pressure (G20 α_{CO})
e_P_DE_G20	$k_B \text{ K cm}^{-3}$	Error on ISM dynamical equilibrium pressure (G20 α_{CO})

(This table is available in its entirety in machine-readable form.)

Σ_{mol} , and P_{DE} are measured on a strictly fixed spatial scale of 1.5 kpc.

To make the improved data products available to the community, we distill all measurements used in this work into a machine-readable table (see Table A1) and publish it with this paper. We will also release an associated new version (v4.0) of the full PHANGS high-level database at the same online location⁴⁸ as the version published in S22.

Appendix B

Fitting SF Laws with `linmix`

In Section 3, we present the best-fit power-law models for each of the four SF laws measured with each of the six methodological

choices. This appendix describes how we determine the best-fit models, the assumptions involved in the model fit, and (most importantly) our treatments for nondetections.

For each SF law measured with each specific set of methodological choice, we determine a best-fit linear relation in the log–log space, which translates to a power-law relation in linear space. We perform this model fit with the `linmix` package (Kelly 2007). The model fit assumes that the two-dimensional data distribution in logarithmic space can be described by a single, underlying linear relation, with residual scatter attributed to a combination of measurement uncertainties (in both x and y) and a fixed intrinsic scatter (along the y direction). The posterior distributions for all model parameters, including nuisance parameters (e.g., those in a Gaussian mixture model describing the data distribution along the x direction), are evaluated in a hierarchical Bayesian framework

⁴⁸ https://www.canfar.net/storage/vault/list/phangs/RELEASES/Sun_etal_2022

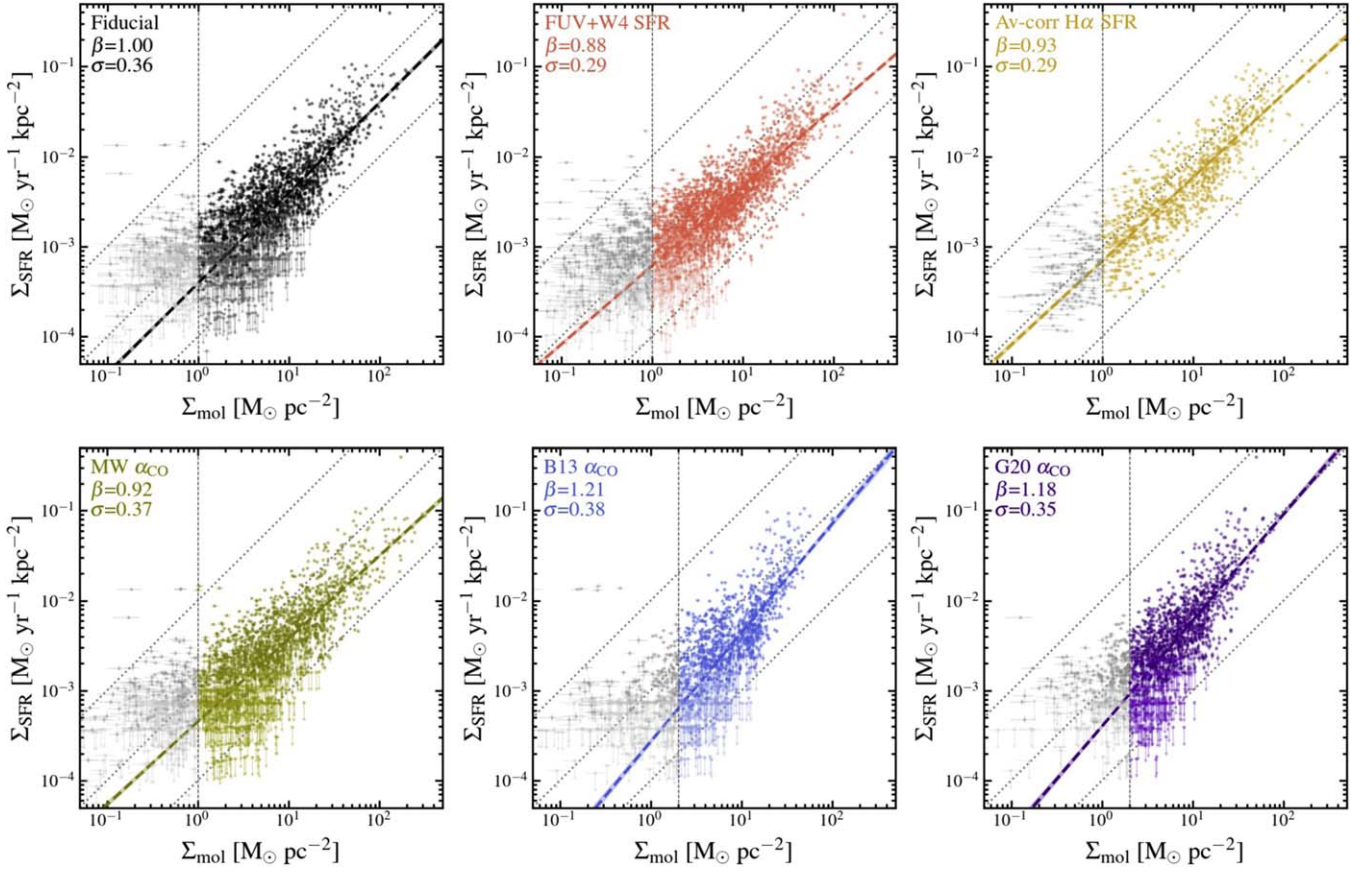


Figure B1. An example figure from the figure set, showing our measurements and best-fit power-law models for the mKS relation and how they vary with methodological choices (i.e., SFR calibrations and CO-to-H₂ conversion factors). In each panel, the best-fit power-law slope β and intrinsic scatter σ are displayed at the top left corner; the vertical dashed line marks the Σ_{mol} threshold above which the power-law model fit is performed ($\sim 1 M_{\odot} \text{pc}^{-2}$ with the fiducial and MW α_{CO} and $\sim 2 M_{\odot} \text{pc}^{-2}$ with the other α_{CO}).

(The complete figure set (four images) is available.)

with Markov Chain Monte Carlo. The best-fit value for each parameter is then determined from the median value over all realizations (see Figure Set B1 for visualizations of the best-fit models). Since the joint posterior distribution of all parameters very closely resembles an N-dimensional Gaussian, using its maxima would yield almost identical results.

The `linmix` model can self-consistently handle data censoring (i.e., nondetections) for the dependent variable but *not* for the independent variable. However, failure to account for the latter could introduce biases due to increasingly incomplete sampling toward the lower end of the relation. This is particularly important for the mKS relation, where the independent variable (Σ_{mol}) is translated directly from CO line intensity and thus strongly affected by a data sensitivity limit at low Σ_{mol} . To address this issue, we calculate the data number density per log Σ_{mol} interval in each Σ_{mol} bin and determine a Σ_{mol} threshold below which the data density drops to near half of the maximum value. This Σ_{mol} threshold roughly coincides with the highest Σ_{mol} upper limits in our sample, which confirms that it is where data censoring becomes important. We thus perform the `linmix` model fit only on measurements (including detections of upper limits for Σ_{SFR}) above this threshold, so that the best-fit parameters can be relatively unaffected. The Σ_{mol} threshold changes according to the adopted CO-to-H₂ conversion factors (see Figure Set B1 for the

exact location of this threshold). This Σ_{mol} clipping is applied to the model fits for all four SF laws examined in Section 3 to ensure self-consistency in this work.

We emphasize that correct treatments of data censoring are crucial for deriving unbiased fit parameters, as has been shown in previous studies (e.g., Pessa et al. 2021). This is especially important when dealing with spatially resolved measurements, for which the number of nondetections can be large and whose distribution is concentrated toward the lower end of the probed parameter space. To illustrate this issue, we perform a test fit for the mKS relation with the fiducial SFR calibration and α_{CO} , but without including any Σ_{SFR} upper limits. This yields a much shallower slope of $\beta = 0.80$ compared with $\beta = 1.00$ when including the upper limits. This is expected since most Σ_{SFR} upper limits are distributed near the low Σ_{mol} end of the mKS relation, without which the average Σ_{SFR} value is biased high at low Σ_{mol} , and thus the power-law slope is biased low. Beside the treatments of data censoring, the handling of measurement uncertainties and choice of regression methods could also affect the fit result (e.g., de los Reyes & Kennicutt 2019; Tabatabaei et al. 2022).

ORCID iDs

Jiayi Sun (孙嘉懿) <https://orcid.org/0000-0003-0378-4667>
Adam K. Leroy <https://orcid.org/0000-0002-2545-1700>

Eve C. Ostriker  <https://orcid.org/0000-0002-0509-9113>
 Sharon Meidt  <https://orcid.org/0000-0002-6118-4048>
 Erik Rosolowsky  <https://orcid.org/0000-0002-5204-2259>
 Eva Schinnerer  <https://orcid.org/0000-0002-3933-7677>
 Christine D. Wilson  <https://orcid.org/0000-0001-5817-0991>
 Dyas Utomo  <https://orcid.org/0000-0003-4161-2639>
 Francesco Belfiore  <https://orcid.org/0000-0002-2545-5752>
 Guillermo A. Blanc  <https://orcid.org/0000-0003-4218-3944>
 Eric Emsellem  <https://orcid.org/0000-0002-6155-7166>
 Christopher Faesi  <https://orcid.org/0000-0001-5310-467X>
 Brent Groves  <https://orcid.org/0000-0002-9768-0246>
 Annie Hughes  <https://orcid.org/0000-0002-9181-1161>
 Eric W. Koch  <https://orcid.org/0000-0001-9605-780X>
 Kathryn Kreckel  <https://orcid.org/0000-0001-6551-3091>
 Daizhong Liu  <https://orcid.org/0000-0001-9773-7479>
 Hsi-An Pan  <https://orcid.org/0000-0002-1370-6964>
 Jérôme Pety  <https://orcid.org/0000-0003-3061-6546>
 Miguel Querejeta  <https://orcid.org/0000-0002-0472-1011>
 Alessandro Razza  <https://orcid.org/0000-0001-7876-1713>
 Toshiaki Saito  <https://orcid.org/0000-0001-9016-2641>
 Amy Sardone  <https://orcid.org/0000-0002-5783-145X>
 Antonio Usero  <https://orcid.org/0000-0003-1242-505X>
 Thomas G. Williams  <https://orcid.org/0000-0002-0012-2142>
 Frank Bigiel  <https://orcid.org/0000-0003-0166-9745>
 Alberto D. Bolatto  <https://orcid.org/0000-0002-5480-5686>
 Mélanie Chevance  <https://orcid.org/0000-0002-5635-5180>
 Daniel A. Dale  <https://orcid.org/0000-0002-5782-9093>
 Jindra Gensior  <https://orcid.org/0000-0001-6119-9883>
 Simon C. O. Glover  <https://orcid.org/0000-0001-6708-1317>
 Kathryn Grasha  <https://orcid.org/0000-0002-3247-5321>
 Jonathan D. Henshaw  <https://orcid.org/0000-0001-9656-7682>
 María J. Jiménez-Donaire  <https://orcid.org/0000-0002-9165-8080>
 Ralf S. Klessen  <https://orcid.org/0000-0002-0560-3172>
 J. M. Diederik Kruijssen  <https://orcid.org/0000-0002-8804-0212>
 Eric J. Murphy  <https://orcid.org/0000-0001-7089-7325>
 Lukas Neumann  <https://orcid.org/0000-0001-9793-6400>
 Yu-Hsuan Teng  <https://orcid.org/0000-0003-4209-1599>
 David A. Thilker  <https://orcid.org/0000-0002-8528-7340>

References

- Astropy Collaboration, Price-Whelan, A. M., Sipőcz, B. M., et al. 2018, *AJ*, 156, 123
- Astropy Collaboration, Robitaille, T. P., Tollerud, E. J., et al. 2013, *A&A*, 558, A33
- Bacchini, C., Fraternali, F., Iorio, G., & Pezzulli, G. 2019, *A&A*, 622, A64
- Barnes, A. T., Longmore, S. N., Battersby, C., et al. 2017, *MNRAS*, 469, 2263
- Barrera-Ballesteros, J. K., Sánchez, S. F., Heckman, T., et al. 2021, *MNRAS*, 503, 3643
- Belfiore, F., Leroy, A. K., Sun, J., et al. 2023, *A&A*, 670, A67
- Belfiore, F., Santoro, F., Groves, B., et al. 2022, *A&A*, 659, A26
- Bigiel, F., Leroy, A., Walter, F., et al. 2008, *AJ*, 136, 2846
- Bigiel, F., Leroy, A. K., Walter, F., et al. 2011, *ApJL*, 730, L13
- Bolatto, A. D., Wolfire, M., & Leroy, A. K. 2013, *ARA&A*, 51, 207
- Calzetti, D., Kennicutt, R. C., Engelbracht, C. W., et al. 2007, *ApJ*, 666, 870
- Chung, A., van Gorkom, J. H., Kenney, J. D. P., Crowl, H., & Vollmer, B. 2009, *AJ*, 138, 1741
- Daddi, E., Elbaz, D., Walter, F., et al. 2010, *ApJL*, 714, L118
- de los Reyes, M. A. C., & Kennicutt, R. C. J. 2019, *ApJ*, 872, 16
- den Brok, J. S., Chatzigiannakis, D., Bigiel, F., et al. 2021, *MNRAS*, 504, 3221
- Elmegreen, B. G. 1997, *RMxAC*, 6, 165
- Emsellem, E., Schinnerer, E., Santoro, F., et al. 2022, *A&A*, 659, A191
- Evans, N. J., Kim, J.-G., & Ostriker, E. C. 2022, *ApJL*, 929, L18
- Federrath, C., & Klessen, R. S. 2012, *ApJ*, 761, 156
- Fisher, D. B., Bolatto, A. D., Glazebrook, K., et al. 2022, *ApJ*, 928, 169
- Fisher, D. B., Bolatto, A. D., White, H., et al. 2019, *ApJ*, 870, 46
- Gao, Y., & Solomon, P. M. 2004, *ApJ*, 606, 271
- Gentry, E. S., Krumholz, M. R., Madau, P., & Lupi, A. 2019, *MNRAS*, 483, 3647
- Genzel, R., Tacconi, L. J., Gracia-Carpio, J., et al. 2010, *MNRAS*, 407, 2091
- Girard, M., Fisher, D. B., Bolatto, A. D., et al. 2021, *ApJ*, 909, 12
- Gong, M., Ostriker, E. C., Kim, C.-G., & Kim, J.-G. 2020, *ApJ*, 903, 142
- Harris, C. R., Millman, K. J., van der Walt, S. J., et al. 2020, *Natur*, 585, 357
- Heald, G., Józsa, G., Serra, P., et al. 2011, *A&A*, 526, A118
- Hennebelle, P., & Chabrier, G. 2011, *ApJL*, 743, L29
- Hunter, J. D. 2007, *CSE*, 9, 90
- Jeffreson, S. M. R., Sun, J., & Wilson, C. D. 2022, *MNRAS*, 515, 1663
- Kado-Fong, E., Kim, C.-G., Greene, J. E., & Lancaster, L. 2022, *ApJ*, 939, 101
- Keller, B. W., Kruijssen, J. M. D., & Chevance, M. 2022, *MNRAS*, 514, 5355
- Kelly, B. C. 2007, *ApJ*, 665, 1489
- Kennicutt, R. C., & Evans, N. J. 2012, *ARA&A*, 50, 531
- Kennicutt, R. C. J. 1989, *ApJ*, 344, 685
- Kennicutt, R. C. J. 1998, *ApJ*, 498, 541
- Kennicutt, R. C. J., & De Los Reyes, M. A. C. 2021, *ApJ*, 908, 61
- Kim, J.-G., Ostriker, E. C., & Filippova, N. 2021, *ApJ*, 911, 128
- Kim, C.-G., Ostriker, E. C., & Kim, W.-T. 2013, *ApJ*, 776, 1
- Kim, C.-G., Ostriker, E. C., & Raileanu, R. 2017, *ApJ*, 834, 25
- Koribalski, B. S., Wang, J., Kamphuis, P., et al. 2018, *MNRAS*, 478, 1611
- Kreckel, K., Faesi, C., Kruijssen, J. M. D., et al. 2018, *ApJL*, 863, L21
- Kregel, M., van der Kruit, P. C., & de Grijs, R. 2002, *MNRAS*, 334, 646
- Krumholz, M. R., Burkhardt, B., Forbes, J. C., & Crocker, R. M. 2018, *MNRAS*, 477, 2716
- Krumholz, M. R., Dekel, A., & McKee, C. F. 2012, *ApJ*, 745, 69
- Krumholz, M. R., & McKee, C. F. 2005, *ApJ*, 630, 250
- Krumholz, M. R., McKee, C. F., & Bland-Hawthorn, J. 2019, *ARA&A*, 57, 227
- Krumholz, M. R., McKee, C. F., & Tumlinson, J. 2009, *ApJ*, 699, 850
- Lada, C. J., Forbrich, J., Lombardi, M., & Alves, J. F. 2012, *ApJ*, 745, 190
- Lang, P., Meidt, S. E., Rosolowsky, E., et al. 2020, *ApJ*, 897, 122
- Leroy, A. K., Bolatto, A. D., Sandstrom, K., et al. 2023, *ApJL*, 944, L10
- Leroy, A. K., Hughes, A., Liu, D., et al. 2021b, *ApJS*, 255, 19
- Leroy, A. K., Rosolowsky, E., Usero, A., et al. 2022, *ApJ*, 927, 149
- Leroy, A. K., Sandstrom, K. M., Lang, D., et al. 2019, *ApJS*, 244, 24
- Leroy, A. K., Schinnerer, E., Hughes, A., et al. 2017, *ApJ*, 846, 71
- Leroy, A. K., Schinnerer, E., Hughes, A., et al. 2021a, *ApJS*, 257, 43
- Leroy, A. K., Walter, F., Brinks, E., et al. 2008, *AJ*, 136, 2782
- Leroy, A. K., Walter, F., Sandstrom, K., et al. 2013, *AJ*, 146, 19
- McKee, C. F., & Ostriker, E. C. 2007, *ARA&A*, 45, 565
- Ostriker, E. C., & Kim, C.-G. 2022, *ApJ*, 936, 137
- Ostriker, E. C., McKee, C. F., & Leroy, A. K. 2010, *ApJ*, 721, 975
- Ostriker, E. C., & Shetty, R. 2011, *ApJ*, 731, 41
- Padoan, P., Federrath, C., Chabrier, G., et al. 2014, in *Protostars and Planets VI, The Star Formation Rate of Molecular Clouds*, ed. H. Beuther et al. (Tucson, AZ: Univ. of Arizona Press), 77
- Padoan, P., Haugbølle, T., & Nordlund, A. 2012, *ApJL*, 759, L27
- Pessa, I., Schinnerer, E., Belfiore, F., et al. 2021, *A&A*, 650, A134
- Querejeta, M., Schinnerer, E., Meidt, S., et al. 2021, *A&A*, 656, A133
- Saintonge, A., & Catinella, B. 2022, *ARA&A*, 60, 319
- Saintonge, A., Kauffmann, G., Wang, J., et al. 2011, *MNRAS*, 415, 61
- Schmidt, M. 1959, *ApJ*, 129, 243
- Schruba, A., Kruijssen, J. M. D., & Leroy, A. K. 2019, *ApJ*, 883, 2
- Silk, J. 1997, *ApJ*, 481, 703
- Sun, J. 2022, PhangsTeam/MegaTable: PHANGS mega-table pipeline public release, v3.0, Zenodo, doi:10.5281/zenodo.6584842
- Sun, J., Leroy, A. K., Ostriker, E. C., et al. 2020a, *ApJ*, 892, 148
- Sun, J., Leroy, A. K., Rosolowsky, E., et al. 2022, *AJ*, 164, 43
- Sun, J., Leroy, A. K., Schinnerer, E., et al. 2020b, *ApJL*, 901, L8
- Sun, J., Leroy, A. K., Schruba, A., et al. 2018, *ApJ*, 860, 172
- Tabatabaei, F. S., Cotton, W., Schinnerer, E., et al. 2022, *MNRAS*, 517, 2990
- Tan, J. C. 2000, *ApJ*, 536, 173
- Thompson, T. A., Quataert, E., & Murray, N. 2005, *ApJ*, 630, 167
- Utomo, D., Sun, J., Leroy, A. K., et al. 2018, *ApJL*, 861, L18
- van der Hulst, J. M., van Albada, T. S., & Sancisi, R. 2001, in *ASP Conf. Ser. 240, Gas and Galaxy Evolution*, ed. J. E. Hibbard, M. Rupen, & J. H. van Gorkom (San Francisco, CA: ASP), 451
- Walter, F., Brinks, E., de Blok, W. J. G., et al. 2008, *AJ*, 136, 2563
- Wilson, C. D., Elmegreen, B. G., Bemis, A., & Brunetti, N. 2019, *ApJ*, 882, 5
- Wong, T., & Blitz, L. 2002, *ApJ*, 569, 157
- Yajima, Y., Sorai, K., Miyamoto, Y., et al. 2021, *PASJ*, 73, 257